

Elevation values for the near side hemisphere
binned in increments of 100 meters. Count is
number of elevation points within the bin. Percent
is count for indicated elevation values
divided by total count.

Minimum	Maximum	Count	Percent
-10000	-9500		
-9500	-9000		
-9000	-8500		
-8500	-8000		
-8000	-7500		
-7500	-7000	31	0.00%
-7000	-6500	304	0.00%
-6500	-6000	760	0.00%
-6000	-5500	3,226	0.01%
-5500	-5000	8,512	0.03%
-5000	-4500	23,964	0.09%
-4500	-4000	54,040	0.20%
-4000	-3500	221,596	0.83%
-3500	-3000	297,377	1.13%
-3000	-2500	1,842,080	6.99%
-2500	-2000	5,422,658	20.00%
-2000	-1500	14,893,310	56.36%
-1500	-1000	5,207,304	19.75%
-1000	-500	4,680,234	17.89%
-500	0	2,266,361	8.61%
0	500	1,245,859	4.76%
500	1000	1,019,491	3.87%
1000	1500	539,764	2.05%
1500	2000	291,199	1.10%
2000	2500	143,507	0.54%
2500	3000	62,633	0.24%
3000	3500	26,652	0.10%
3500	4000	10,602	0.04%
4000	4500	4,440	0.02%
4500	5000	2,616	0.01%
5000	5500	1,250	0.00%
5500	6000	559	0.00%
6000	6500	131	0.00%
6500	7000		
7000	7500		
7500	8000		
8000	8500		
8500	9000		
9000	9500		
9500	10000		

Elevation values for the far side hemisphere
binned in increments of 100 meters. Count is
number of elevation points within the bin. Percent
is count for indicated elevation values
divided by total count.

Minimum	Maximum	Count	Percent
-10000	-9500		
-9500	-9000		
-9000	-8500	1,464	0.01%
-8500	-8000	2,878	0.01%
-8000	-7500	4,541	0.02%
-7500	-7000	8,761	0.03%
-7000	-6500	38,139	0.14%
-6500	-6000	52,984	0.20%
-6000	-5500	146,802	0.56%
-5500	-5000	599,121	2.25%
-5000	-4500	1,316,072	4.87%
-4500	-4000	2,773,935	10.23%
-4000	-3500	5,151,749	19.25%
-3500	-3000	11,253,599	42.61%
-3000	-2500	63,644	0.24%
-2500	-2000	1,573,795	5.95%
-2000	-1500	2,176,386	8.22%
-1500	-1000	2,091,301	7.91%
-1000	-500	1,854,489	7.01%
-500	0	1,899,232	7.18%
0	500	1,918,159	7.25%
500	1000	1,316,072	4.87%
1000	1500	1,573,935	5.95%
1500	2000	1,316,072	4.87%
2000	2500	5,151,749	19.25%
2500	3000	11,253,599	42.61%
3000	3500	400	0.00%
3500	4000	856,299	3.24%
4000	4500	761,785	2.88%
4500	5000	619,379	2.34%
5000	5500	443,699	1.68%
5500	6000	318,252	1.20%
6000	6500	238,135	0.90%
6500	7000	80,376	0.30%
7000	7500	23,240	0.09%
7500	8000	6,218	0.02%
8000	8500	1,463	0.01%
8500	9000		
9000	9500		
9500	10000		

Histogram of elevation values for near side hemisphere

Histogram of elevation values for far side hemisphere

Scale 1:10,000,000 (1 cm = 10 km)
Lambert Azimuthal Equal-Area projection
LUNAR NEAR SIDE HEMISPHERE
L 10M 00 RTK

NOTES ON BASE
This sheet is one in a series of topographic maps that presents color-coded topographic data digitally merged with shaded relief data.

ADOPTED FIGURE
The figure for the Moon, used for the computation of the map projection, is a sphere with a radius of 1737.4 km (Staidman and others, 2003). Because the Moon has no surface water, and hence no sea level, the datum (the 0 km contour) for elevations is defined as the radius of 1737.4 km. Coordinates are based on the mean Earth-polar axis (M.E.P.A.) coordinate system, the z axis is the axis of the Moon's rotation, and the x axis is the mean Earth direction. The center of mass is the origin of the coordinate system (Davies and Colvin, 2000). The equator lies in the x-y plane and the prime meridian lies in the x-z plane with east longitude values being positive.

PROJECTION
The projection is Lambert Azimuthal Equal Area Projection. The scale factor at the central latitude and central longitude point is 1:10,000,000. For the near side hemisphere the central latitude and central longitude point is at 0° and 0°. For the far side hemisphere the central latitude and central longitude point is at 0° and 180°.

CONTROL
The original control for the shaded relief maps was based on horizontal control used in the Lunar Positional Reference of 1973 (Schlatterman, 1975) of the Apollo control system of 1973. Positional discrepancies as large as 25 km at any scale existed in the original shaded relief base. To improve the accuracy, digital shaded relief data were aligned with a mosaic produced from Clementine 750 nm images (Elison, 1997; Isbell and others, 1997; Elison and others, 1999). This alignment process consisted of picking features (commonly called match or tie points) that were visible in both the shaded relief data and the Clementine mosaic. To accomplish this, the files were divided into three areas: north pole, equatorial region, and south pole. They were aligned first in the equatorial region and then in the polar regions. Within the equatorial region, an area extending from 60° S to 60° N, approximately 1100 points were picked. Within the north polar region, an area from 57° N to 90° N, approximately 1900 points were picked. Within the south polar region, an area from 57° S to 90° S, approximately 1100 points were picked. These points were used to warp the shaded relief map to match the Clementine mosaic. The Clementine mosaic has a positional accuracy of 500 to 1000 m (Elison and others, 1999). Vertical control is based on measurements from the Clementine laser altimeter that collected data between 79° S and 81° N. The along-track spacing of these measurements varied over time, with the instrument along-track spacing of 20 km was achieved, where the instrument lost lock over rough highland terrain, the spacing degraded to 100 km.

The shaded relief was based on the orbital ground track and is approximately 60 km (27°) at the equator. Elevation values were collected at 72,548 points by the Clementine laser altimeter. The estimated vertical accuracy of points collected by the Clementine laser altimeter is 130 m (Smith and others, 1997). The Clementine laser altimeter did not collect data over the lunar north or south pole. Therefore, topographic data were collected photogrammetrically to fill in these gaps. The image sources were the oblique and nadir 750 nm images collected by the Clementine UVVIS camera (Rosick and others, 1998). Vertical control, for the photogrammetric data, was established by using the Clementine laser altimeter data at the outer edge of these circular polar areas, and the imagery was used to bridge control and fill in the central part of the circle (Rosick and others, 2001). The expected precision of points collected photogrammetrically is 180 m (Rosick and others, 1998). Further discussion of the photogrammetric topographic data can be found in the topographic data section below.

IMAGE BASE
The shaded relief data were originally published as a series of 1.5 million shaded relief maps. This series included three U.S. Geological Survey maps: I-1218-B, Shaded Relief Map of the Lunar Far Side, 1988; I-1326-A, Shaded Relief Map of the Lunar Polar Regions, 1981; and I-2278, Sheet 2 of 2, Shaded Relief Map of the Lunar Near Side, 1992. These data were digitized and mosaicked into a single digital file. An area of approximately 500,000 km² near the south pole was not visible in any pre-Clementine images and is blank on the published map. The digitized relief base was revised based on the Clementine mosaic and recent Earth-based radar imagery (Margot and others, 1999) to show features in this area. Errors that were present in the original interpretations of lunar morphology have been corrected in the digital version of the warped shaded relief base. These original errors were caused by scanty data, ambiguities introduced by highly oblique solar incidence angles, and distortions created in generating orthorectified images from oblique imagery (Rosick and Anselmann, 2001).

TOPOGRAPHIC DATA
The Clementine laser altimeter points were interpolated to create a global topographic gridded digital terrain model for the lunar surface. Because the altimeter points were sparser near the poles and non-existent over the poles in this digital terrain model, only data between 75° S and 75° N were used in the final digital terrain model. To fill in the polar regions, topographic data were collected photogrammetrically from Clementine 750 nm oblique and nadir images. For the photogrammetric analysis, horizontal control was established by selecting some of the match points that were used in building the Clementine mosaic. These points provided estimates for latitude and longitude values, but no estimate for elevation values. Vertical control

was established by using the global topographic gridded digital terrain model developed from the Clementine laser altimeter points to estimate elevation values for the Clementine match points. To improve the geometry of the control network for the photogrammetric analysis, the Clementine match points, which were selected to tie two images together in order to build the Clementine global mosaic, were transferred to all images that contained the point. Additional points were selected to have four well-distributed points per image, where possible. Analytical aerotriangulation, a weighted least squares process, is used to solve for the parameters of the photogrammetric project. These parameters include image sensor position and angles; latitude, longitude, and elevation of match points; and image coordinates of match points. Adjusting the weight assigned to a parameter determines whether values with high weight are held to the original estimate or values with low weight are allowed to float and a new value determined for the parameter. The parameters with the most error in their original estimate for their values are the image sensor angles, so they are given a low weight. The latitude and longitude values of Clementine match points are given a high weight to the solution holds to the Clementine global mosaic horizontal coordinates. Weights for the elevation values are varied depending on the horizontal distance to a Clementine laser altimeter point: match points within 2000 m of a Clementine laser altimeter point are given a high weight; match points between 2000 m and 5000 m from a Clementine laser altimeter point are given a medium weight; match points greater than 5000 m from a Clementine laser altimeter point are given a low weight. This weighting allows the vertical control to be bridged between areas of known vertical control (the area covered by Clementine laser altimeter measurements) and the areas void of control (the area over the poles).

The Clementine mission collected both oblique and vertical images over the polar regions; these images form stereo pairs that can be used photogrammetrically to collect topographic data. In the north pole region (90° N to 64° N latitude) data were collected from 640 stereo models. Over the south pole region (90° S to 63° S latitude) topographic data were collected from 667 stereo models. Topographic data were collected within each stereo model at an elevation point spacing of 1 km in the x and y directions. This spacing resulted in 1,437,368 points being collected in the north pole region and 1,274,872 points being collected in the south pole region. On average, over the area that was collected in the north pole region (79° S and 81° N), the Clementine laser altimeter collected an elevation value for every 514 km², and the photogrammetric data collected an elevation value for every 1.3 km² in the north pole and 1.2 km² in the south pole. The photogrammetric data were merged and vertically transferred to align with the Clementine laser altimeter data to form the final digital terrain model (Rosick and others, 2001). Merging the topographic data required an iterative process to reduce the error between the photogrammetric data and the topographic data

derived from the Clementine laser altimeter data. At first the photogrammetric topographic data exhibited a systematic stair-step error, in that stereo models closer to the poles had a systematic bias to be higher than stereo models further from the poles. When this bias was removed, the resulting digital terrain model had a bowl-shape appearance with a low spot near the pole. Each stereo model was tilted to remove the bowl shape for stereo models in the north pole area this was 800 m per degree and for stereo models in the south pole area this was 1 km per degree. To fit the stereo models to the Clementine laser altimeter data, an initial set of stereo models that overlaid the Clementine laser altimeter data was adjusted to vertically align with the Clementine laser altimeter data. Stereo models that overlaid the previously adjusted stereo models were adjusted further, this process continued until all stereo models were adjusted to align vertically.

In analyzing the photogrammetric data for the north pole region, 1,189,935 points (83% of the data) occur on two or more stereo models. For 90% of those points, the standard deviation in elevation is between 24 and 409 m, with a mean of 184 m. In the south pole region, 1,262,349 points (79% of the data) occur on two or more stereo models. For 90% of those points, the standard deviation in elevation is between 13 and 395 m, with a mean of 160 m.

In the north pole region 1,276 Clementine laser altimeter points are in the area where photogrammetric topographic data were collected. Comparing the photogrammetric topographic data with the Clementine laser altimeter points shows that for the north pole region the stereo topographic data is an average of 46 m higher than the Clementine laser altimeter points, with a standard deviation of 719 m. In the south pole region, 1,379 Clementine laser altimeter points are in the area where photogrammetric topographic data were collected. For the south pole region the stereo topographic data is an average of 163 m lower than the Clementine laser altimeter points, with a standard deviation of 1005 m.

The merged topographic data were color-coded and combined with the shaded relief data. Areas where no topographic data were collected are not color-coded, and the shaded relief image is shown as gray scale.

NOMENCLATURE
The number, size, and placement of text annotations were chosen to provide a general orientation of conspicuous features on a 1:10,000,000-scale map. Features are labeled with names approved by the International Astronomical Union (for a complete list of lunar nomenclature, see <http://planetarynames.usgs.gov/>).

L 10M 00 RTK: Abbreviation for Moon; 1:10,000,000 series; center of map; 0° latitude and 0° longitude for near side hemisphere, and 0° latitude and 180° longitude for far side hemisphere; shaded relief (R) with topographic data and

OTHER SOURCES OF LUNAR ELEVATION DATA
Two other recent sources of lunar elevation data are from work done by Dr. Tony Cook of the National Air and Space Museum and by Dr. Jean-Luc Margot while at Cornell University. The elevation data collected by Dr. Cook is derived from overlapping Clementine images from the same orbit. This method provides a larger set of images to use in collecting topographic data. Elevation values that are collected are relative to each other and are not known in an absolute sense relative to some datum. To provide an estimate of the absolute height, the elevation values are matched to the Clementine altimeter data and to elevation values previously collected where Clementine altimeter data does not exist. Traditional photogrammetric solutions employ a least squares method to distribute the error throughout the model, whereas the method employed by Dr. Cook aligns the derived data with the Clementine altimeter data where it exists and in areas where there is no data the relative elevation measurements are aligned with previous aligned data until the area is connected back with Clementine data. The error in unknown areas is unconstrained, but the method does close back to an area with known elevation values, which constrains the error at the boundaries of the data (Cook and others, 2000).

The elevation data collected by Dr. Margot is derived from radar interferometry using the Goldstone radar. Over the polar region, five images were collected and mosaicked together. To collect elevation information, the radar signal is sent from one antenna and is received by a least two antennas. The phase difference in the received signal from the two radar sites provides information about the elevation. The elevation information derived from radar interferometry is relative and requires the Clementine altimeter data to provide absolute elevation information relative to a datum (Margot and others, 1999, 2000).

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Scale 1:10,000,000 (1 cm = 10 km)
Lambert Azimuthal Equal-Area projection
LUNAR FAR SIDE HEMISPHERE
L 10M 0180 RTK

NOTES ON BASE
This sheet is one in a series of topographic maps that presents color-coded topographic data digitally merged with shaded relief data.

ADOPTED FIGURE
The figure for the Moon, used for the computation of the map projection, is a sphere with a radius of 1737.4 km (Staidman and others, 2003). Because the Moon has no surface water, and hence no sea level, the datum (the 0 km contour) for elevations is defined as the radius of 1737.4 km. Coordinates are based on the mean Earth-polar axis (M.E.P.A.) coordinate system, the z axis is the axis of the Moon's rotation, and the x axis is the mean Earth direction. The center of mass is the origin of the coordinate system (Davies and Colvin, 2000). The equator lies in the x-y plane and the prime meridian lies in the x-z plane with east longitude values being positive.

PROJECTION
The projection is Lambert Azimuthal Equal Area Projection. The scale factor at the central latitude and central longitude point is 1:10,000,000. For the near side hemisphere the central latitude and central longitude point is at 0° and 0°. For the far side hemisphere the central latitude and central longitude point is at 0° and 180°.

CONTROL
The original control for the shaded relief maps was based on horizontal control used in the Lunar Positional Reference of 1973 (Schlatterman, 1975) of the Apollo control system of 1973. Positional discrepancies as large as 25 km at any scale existed in the original shaded relief base. To improve the accuracy, digital shaded relief data were aligned with a mosaic produced from Clementine 750 nm images (Elison, 1997; Isbell and others, 1997; Elison and others, 1999). This alignment process consisted of picking features (commonly called match or tie points) that were visible in both the shaded relief data and the Clementine mosaic. To accomplish this, the files were divided into three areas: north pole, equatorial region, and south pole. They were aligned first in the equatorial region and then in the polar regions. Within the equatorial region, an area extending from 60° S to 60° N, approximately 1100 points were picked. Within the north polar region, an area from 57° N to 90° N, approximately 1900 points were picked. Within the south polar region, an area from 57° S to 90° S, approximately 1100 points were picked. These points were used to warp the shaded relief map to match the Clementine mosaic. The Clementine mosaic has a positional accuracy of 500 to 1000 m (Elison and others, 1999). Vertical control is based on measurements from the Clementine laser altimeter that collected data between 79° S and 81° N. The along-track spacing of these measurements varied over time, with the instrument along-track spacing of 20 km was achieved, where the instrument lost lock over rough highland terrain, the spacing degraded to 100 km.

Scale 1:10,000,000 (1 cm = 10 km)
Lambert Azimuthal Equal-Area projection
LUNAR FAR SIDE HEMISPHERE
L 10M 0180 RTK

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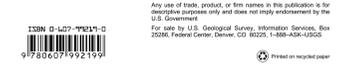
PROJECTION
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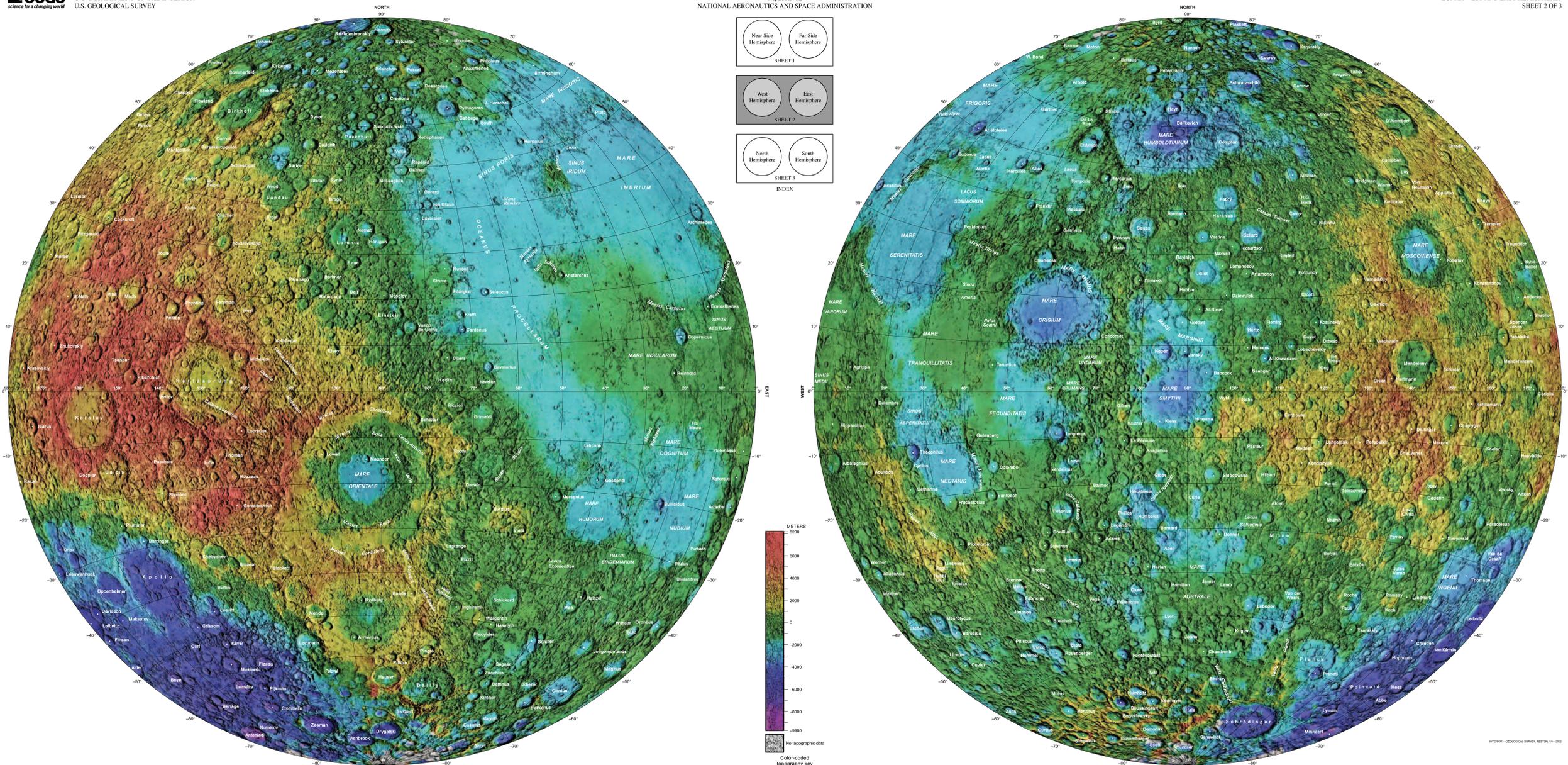
CONTROL
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Histogram of elevation values for near side hemisphere

Histogram of elevation values for far side hemisphere

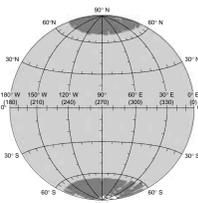
COLOR-CODED TOPOGRAPHY AND SHADED RELIEF MAP OF THE LUNAR NEAR SIDE AND FAR SIDE HEMISPHERES
L 10M 0/0 180 RTK
2002





Elevation values for the west hemisphere binned in increments of 500 meters. Count is number of elevation points within the bin. Percent is count for indicated elevation values divided by total count.

Minimum	Maximum	Count	Percent
-1000	-950	169	0.00%
-950	-900	518	0.00%
-900	-850	826	0.00%
-850	-800	2,380	0.01%
-800	-750	4,147	0.02%
-750	-700	7,245	0.03%
-700	-650	35,531	0.13%
-650	-600	48,612	0.18%
-600	-550	110,030	0.42%
-550	-500	348,543	1.22%
-500	-450	336,099	1.20%
-450	-400	380,951	1.44%
-400	-350	297,589	1.13%
-350	-300	277,557	1.05%
-300	-250	998,766	3.78%
-250	-200	2,428,094	9.29%
-200	-150	5,313,526	11.87%
-150	-100	2,818,781	10.69%
-100	-50	2,223,878	8.33%
-50	0	1,744,190	6.61%
0	50	1,491,147	5.61%
50	100	1,428,908	5.41%
100	150	1,138,869	4.32%
150	200	1,074,097	3.92%
200	250	1,072,031	4.06%
250	300	1,062,262	4.03%
300	350	826,708	3.13%
350	400	701,355	2.66%
400	450	737,026	2.74%
450	500	662,093	2.39%
500	550	441,796	1.67%
550	600	317,694	1.20%
600	650	237,041	0.90%
650	700	20,408	0.30%
700	750	23,217	0.09%
750	800	6,233	0.02%
800	850	1,459	0.01%
850	900		
900	950		
950	1000		



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ADOPTED FIGURE
The figure for the Moon, used for the computation of the map projection, is a sphere with a radius of 1737.4 km (Staidman and others, 2003). Because the Moon has no surface water, and hence no sea level, the datum (the 0 km contour) for elevations is defined as the radius of 1737.4 km. Coordinates are based on the mean Earth-polar axis (M.E.C.) coordinate system, the z axis is the axis of the Moon's rotation, and the x axis is the mean Earth direction. The center of mass is the origin of the coordinate system (Davies and Colvin, 2000). The equator lies in the x-y plane and the prime meridian lies in the x-z plane with east longitude values being positive.

PROJECTION
The projection is Lambert Azimuthal Equal Area Projection. The scale factor at the central latitude and central longitude point is 1:10,000,000. For the west side hemisphere the central latitude and central longitude point is 0° and 90° west. For the east side hemisphere the central latitude and central longitude point is 0° and 90° east.

CONTROL
The original control for the shaded relief maps was based on horizontal control tied to either the Lunar Positional Reference of 1973 (Schlumberger, 1975) or the Apollo control system of 1973. Positional discrepancies as large as 25 km at map scale existed in the original shaded relief base. To improve the accuracy, digital shaded relief data were aligned with a mosaic produced from Clementine 750 nm images (Elison, 1997; Isbell and others, 1997; Elison and others, 1999). This alignment process consisted of picking features (commonly called match or tie points) that were visible in both the shaded relief and the Clementine mosaic. To accomplish this, the files were divided into three areas: north pole, equatorial region, and south pole. They were aligned first in the equatorial region and then in the polar regions. Within the equatorial region, an area extending from 60° S. to 60° N., approximately 1100 points were picked. Within the north polar region, an area from 57° N. to 90° N., approximately 1900 points were picked. Within the south polar region, an area from 57° S. to 90° S., approximately 1100 points were picked. These points were used to warp the shaded relief map to match the Clementine mosaic. The Clementine mosaic has a positional accuracy of 90 m (1.0667) (Elison and others, 1999).

Vertical control is based on measurements from the Clementine laser altimeter that collected data between 79° S. and 81° N. The along-track spacing of these measurements varied, over some smooth surface an along-track spacing of 20 km was achieved, where the instrument lost lock over rough highland terrain, the spacing degraded to 100 km.

The across-track spacing was based on the orbital ground track and is approximately 60 km (2) at the equator. Elevation values were collected at 72,548 points by the Clementine laser altimeter. The estimated vertical accuracy of points collected by the Clementine laser altimeter is 130 m (Smith and others, 1997). The Clementine laser altimeter did not collect data over the lunar north or south pole. Therefore, topographic data were collected photogrammetrically to fill in these gaps. The image sources were the oblique and nadir 750 nm images collected by the Clementine UVVIS camera (Rosick and others, 1998). Vertical control, for the photogrammetric data, was established by using the Clementine laser altimeter data at the outer edge of these circular polar areas, and the imagery was used to bridge control and fill in the central part of the circle (Rosick and others, 2001). The expected precision of points collected photogrammetrically is 180 m (Rosick and others, 1998). Further discussion of the photogrammetric topographic data can be found in the topographic data section below.

IMAGE BASE
The shaded relief data were originally published as a series of 1.5 million shaded relief maps. This series included three U.S. Geological Survey maps: I-1218-B, Shaded Relief Map of the Lunar Far Side, 1989; I-1326-A, Shaded Relief Map of the Lunar Near Side, 1981; and I-2278, Sheet 2 of 2, Shaded Relief Map of the Lunar Near Side, 1992. These data were digitized and mosaicked into a single digital file. An area of approximately 500,000 km² near the south pole was not visible in any pre-Clementine images and is blank on the published map. The digitized relief base was revised based on the Clementine mosaic and recent Earth-based radar imagery (Margot and others, 1999) to show features in this area. Errors that were present in the original interpretations of lunar morphology have been corrected in the digital version of the warped shaded relief map base. These original errors were caused by scanty data, ambiguities introduced by highly oblique solar incidence angles, and distortions created in generating orthorectified images from oblique images (Rosick and Ansellman, 2001).

TOPOGRAPHIC DATA
The Clementine laser altimeter points were interpolated to create a global topographic gridded digital terrain model for the lunar surface. Because the altimeter points were sparser near the poles and non-existent over the poles in this digital terrain model, only data between 75° S. and 75° N. were used in the final digital terrain model. To fill in the polar regions, topographic data were collected photogrammetrically from Clementine 750 nm oblique and nadir images. For the photogrammetric analysis, horizontal control was established by selecting some of the match points that were used in building the Clementine global mosaic. These points provided estimates for latitude and longitude values, but no estimate for elevation values. Vertical control

was established by using the global topographic gridded digital terrain model developed from the Clementine laser altimeter points to estimate elevation values for the Clementine match points. To improve the geometry of the control network for the photogrammetric analysis, the Clementine match points, which were selected to tie two images together in order to build the Clementine global mosaic, were transferred to all images that contained the point. Additional points were selected to have four well-distributed points per image, where possible. Analytical aerotriangulation, a weighted least squares process, is used to solve for the parameters of the photogrammetric project. These parameters include image sensor position and angles; latitude, longitude, and elevation of match points; and image coordinates of match points. Adjusting the weight assigned to a parameter determines whether values with high weight are held to the original estimate or values with low weight are allowed to float and a new value determined for the parameter. The parameters with the most error in their original estimate for their values are the image sensor angles, so they are given a low weight. The latitude and longitude values of Clementine match points are given a high weight to the solution holds to the Clementine global mosaic horizontal coordinates. Weights for the elevation values are varied depending on the horizontal distance to a Clementine laser altimeter point: match points within 2000 m of a Clementine laser altimeter point are given a high weight; match points between 2000 m and 5000 m from a Clementine laser altimeter point are given a medium weight; and match points greater than 5000 m from a Clementine laser altimeter point are given a low weight. This weighting allows the vertical control to be bridged between areas of known vertical control (the area covered by Clementine laser altimeter measurements) and the areas void of control (the area over the poles).

The Clementine mission collected both oblique and vertical images over the polar regions; these images from stereo pairs that can be used photogrammetrically to collect topographic data. In the north pole region (90° N. to 64° N. latitude) data were collected from 640 stereo models. Over the south pole region (90° S. to 63° S. latitude) topographic data were collected from 667 stereo models. Topographic data were collected within each stereo model with an elevation point spacing of 1 km in the x and y directions. This spacing resulted in 1,473,368 points being collected in the north pole region and 1,724,872 points being collected in the south pole region. On average, over the area that data were collected (79° S. and 81° N.), the Clementine laser altimeter data collected an elevation value for every 514 km², and the photogrammetric data collected an elevation value for every 1.3 km² in the north pole and 1.2 km² in the south pole. The photogrammetric data were merged and vertically transferred to align with the Clementine laser altimeter data to form the final digital terrain model (Rosick and others, 2001). Merging the topographic data required an iterative process to reduce the error between the photogrammetric data and the topographic data

derived from the Clementine laser altimeter data. At first the photogrammetric topographic data exhibited a systematic stair-step error, in that stereo models closer to the poles had a systematic bias to be higher than stereo models farther from the poles. When this bias was removed, the resulting digital terrain model had a bowl-shape appearance with a low spot near the pole. Each stereo model was tilted to remove the bowl shape for stereo models in the north pole area this was 800 m per degree and for stereo models in the south pole area this was 1 km per degree. To fit the stereo models to the Clementine laser altimeter data, an initial set of stereo models that overlaid the Clementine laser altimeter data was adjusted to vertically align with the Clementine laser altimeter data. Stereo models that overlaid the previously adjusted stereo models were adjusted further, this process continued until all stereo models were adjusted to align vertically. In analyzing the photogrammetric data for the north pole region, 1,189,935 points (81% of the data) occur on two or more stereo models. For 90% of those points, the standard deviation in elevation is between 24 and 409 m, with a mean of 184 m. In the south pole region, 1,262,349 points (73% of the data) occur on two or more stereo models. For 90% of those points, the standard deviation in elevation is between 13 and 395 m, with a mean of 160 m. In the north pole region 1,276 Clementine laser altimeter points are in the area where photogrammetric topographic data were collected. Comparing the photogrammetric topographic data with the Clementine laser altimeter points shows that for the north pole region the stereo topographic data is an average of 46 m higher than the Clementine laser altimeter points, with a standard deviation of 719 m. In the south pole region, 1,379 Clementine laser altimeter points are in the area where photogrammetric topographic data were collected. For the south pole region the stereo topographic data is an average of 163 m lower than the Clementine laser altimeter points, with a standard deviation of 1005 m.

The merged topographic data were color-coded and combined with the shaded relief data. Areas where no topographic data were collected are not color-coded, and the shaded relief image is shown as gray scale.

NOMENCLATURE
The number, size, and placement of text annotations were chosen to provide a general orientation of conspicuous features on a 1:10,000,000-scale map. Features are labeled with names approved by the International Astronomical Union (for a complete list of lunar nomenclature, see <http://planetarynames.wr.usgs.gov>). L 10M 0/450 RTK: Abbreviation for Moon; 1:10,000,000 series; center of maps: 0° latitude and 90° west longitude for west hemisphere, and 0° latitude and 90° east longitude for east hemisphere; shaded relief (R) with topographic data and nomenclature (T) and

color-coded topographic data (K) (Greely and Basov, 1990).

OTHER SOURCES OF LUNAR ELEVATION DATA
Two other recent sources of lunar elevation data are from work done by Dr. Tony Cook of the National Air and Space Museum and by Dr. Jean-Luc Margot while at Cornell University. The elevation data collected by Dr. Cook is derived from overlapping Clementine images from the same orbit. This method provides a larger set of images to use in collecting topographic data. Elevation values that are collected are relative to each other and are not known in an absolute sense relative to some datum. To provide an estimate of the absolute height, the elevation values are matched to the Clementine altimeter data and to elevation values previously collected where Clementine altimeter does not exist. Traditional photogrammetric solutions employ a least squares method to distribute the error throughout the model, whereas the method employed by Dr. Cook aligns the derived data with the Clementine altimeter data where it exists and in areas where there is no data the relative elevation measurements are aligned with previous aligned data until the area is connected back with Clementine data. The error in unknown areas is unconstrained, but the method does close back to an area with known elevation values, which constrains the error at the starting and ending spots (Cook and others 2000).

The elevation data collected by Dr. Margot is derived from radar interferometry using the Goldstone radar. Over the polar region, five images were collected and mosaicked together. To collect elevation information, the radar signal is sent from one antenna and is received by a least two antennas. The phase difference in the received signal from the two radar sites provides information about the elevation. The elevation information derived from radar interferometry is relative and requires the use of the Clementine altimeter points to provide absolute elevation information relative to a datum (Margot and others 1999, 2000).

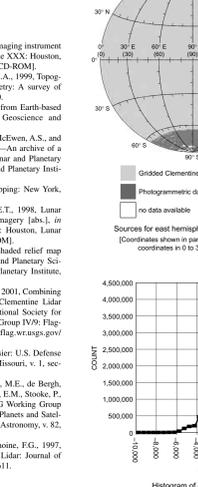
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Elevation values for the east hemisphere binned in increments of 500 meters. Count is number of elevation points within the bin. Percent is count for indicated elevation values divided by total count.

Minimum	Maximum	Count	Percent
-1000	-950	119	0.00%
-950	-900	238	0.00%
-900	-850	635	0.00%
-850	-800	473	0.00%
-800	-750	312	0.00%
-750	-700	1,054	0.00%
-700	-650	2,013	0.01%
-650	-600	5,157	0.02%
-600	-550	40,190	0.15%
-550	-500	179,536	0.68%
-500	-450	201,306	0.76%
-450	-400	472,507	1.79%
-400	-350	659,666	2.50%
-350	-300	1,507,604	5.71%
-300	-250	1,176,181	4.50%
-250	-200	2,782,187	10.54%
-200	-150	3,956,635	14.99%
-150	-100	4,027,048	15.20%
-100	-50	3,006,987	11.39%
-50	0	1,907,225	7.23%
0	50	1,489,511	5.64%
50	100	1,139,180	5.00%
100	150	1,073,012	4.07%
150	200	845,010	3.20%
200	250	523,506	1.98%
250	300	325,416	1.23%
300	350	400,167	1.63%
350	400	43,277	0.16%
400	450	16,225	0.06%
450	500	3,231	0.01%
500	550	1,113	0.00%
550	600	324	0.00%
600	650		
650	700		
700	750		
750	800		
800	850		
850	900		
900	950		
950	1000		

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NOTES ON BASE
This sheet is one in a series of topographic maps that presents color-coded topographic data digitally merged with shaded relief data.

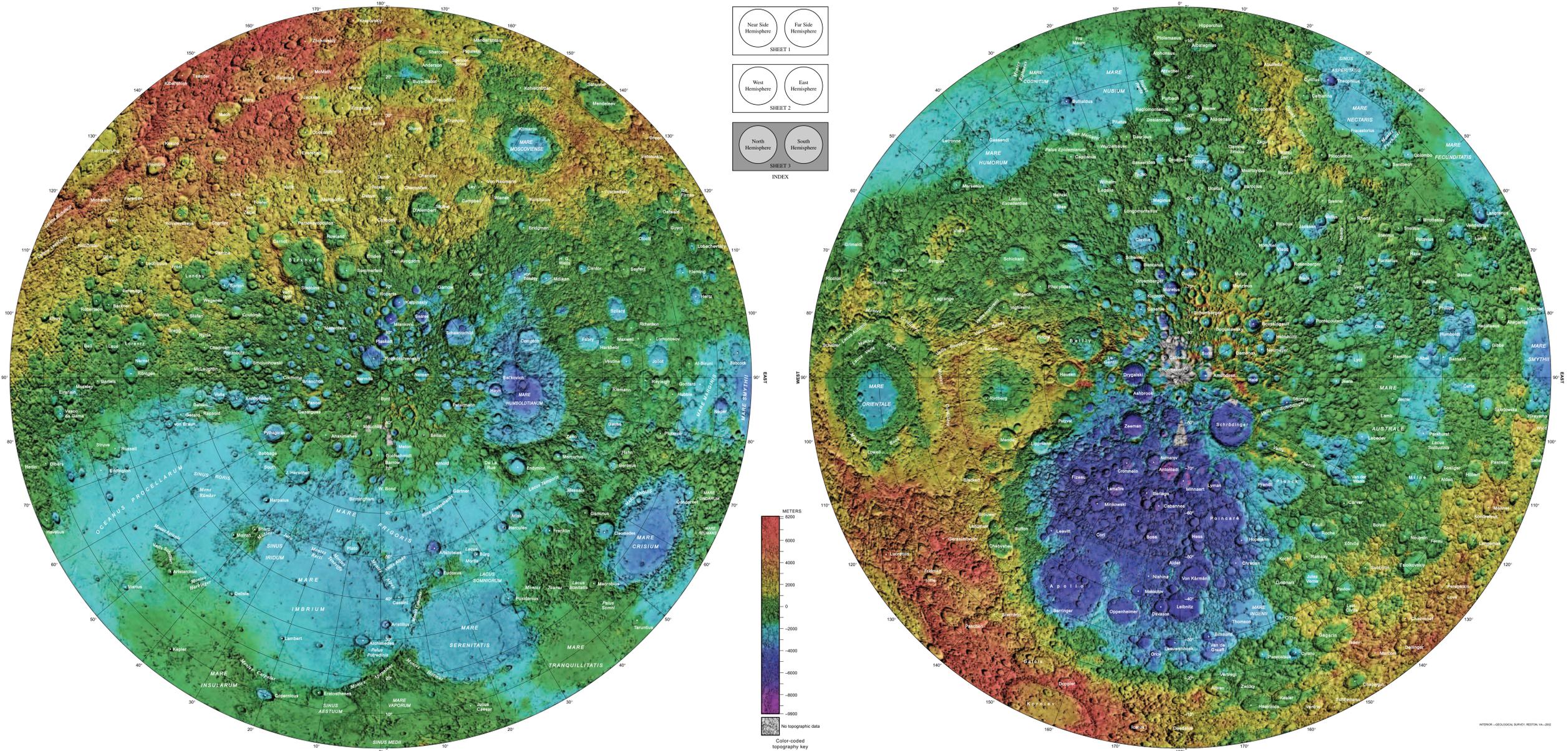
ADOPTED FIGURE
The figure for the Moon, used for the computation of the map projection, is a sphere with a radius of 1737.4 km (Staidman and others, 2003). Because the Moon has no surface water, and hence no sea level, the datum (the 0 km contour) for elevations is defined as the radius of 1737.4 km. Coordinates are based on the mean Earth-polar axis (M.E.C.) coordinate system, the z axis is the axis of the Moon's rotation, and the x axis is the mean Earth direction. The center of mass is the origin of the coordinate system (Davies and Colvin, 2000). The equator lies in the x-y plane and the prime meridian lies in the x-z plane with east longitude values being positive.

PROJECTION
The projection is Lambert Azimuthal Equal Area Projection. The scale factor at the central latitude and central longitude point is 1:10,000,000. For the west side hemisphere the central latitude and central longitude point is 0° and 90° west. For the east side hemisphere the central latitude and central longitude point is 0° and 90° east.

CONTROL
The original control for the shaded relief maps was based on horizontal control tied to either the Lunar Positional Reference of 1973 (Schlumberger, 1975) or the Apollo control system of 1973. Positional discrepancies as large as 25 km at map scale existed in the original shaded relief base. To improve the accuracy, digital shaded relief data were aligned with a mosaic produced from Clementine 750 nm images (Elison, 1997; Isbell and others, 1997; Elison and others, 1999). This alignment process consisted of picking features (commonly called match or tie points) that were visible in both the shaded relief and the Clementine mosaic. To accomplish this, the files were divided into three areas: north pole, equatorial region, and south pole. They were aligned first in the equatorial region and then in the polar regions. Within the equatorial region, an area extending from 60° S. to 60° N., approximately 1100 points were picked. Within the north polar region, an area from 57° N. to 90° N., approximately 1900 points were picked. Within the south polar region, an area from 57° S. to 90° S., approximately 1100 points were picked. These points were used to warp the shaded relief map to match the Clementine mosaic. The Clementine mosaic has a positional accuracy of 90 m (1.0667) (Elison and others, 1999).

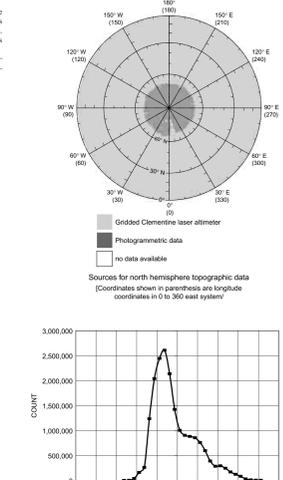
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**COLOR-CODED TOPOGRAPHY AND SHADED RELIEF MAP OF THE LUNAR WEST AND EAST HEMISPHERES
L 10M 0/4-90 RTK
2002**



Elevation values for the north hemisphere binned in increments of 500 meters. Count is number of elevation points within the bin. Percent is count for indicated elevation values divided by total count.

Minimum	Maximum	Count	Percent
-10000	-9500	222	0.00%
-9500	-9000	543	0.01%
-9000	-8500	1,079	0.01%
-8500	-8000	2,074	0.01%
-8000	-7500	3,179	0.02%
-7500	-7000	6,236	0.03%
-7000	-6500	10,393	0.05%
-6500	-6000	18,548	0.09%
-6000	-5500	32,705	0.16%
-5500	-5000	58,862	0.29%
-5000	-4500	104,019	0.52%
-4500	-4000	191,176	0.95%
-4000	-3500	351,333	1.72%
-3500	-3000	650,490	3.17%
-3000	-2500	1,235,289	6.02%
-2500	-2000	2,308,202	11.35%
-2000	-1500	4,244,140	20.59%
-1500	-1000	7,690,187	37.56%
-1000	-500	13,125,259	63.86%
-500	0	14,222,579	69.00%
0	500	10,022,480	48.99%
500	1000	902,322	4.36%
1000	1500	88,066	0.43%
1500	2000	8,427	0.04%
2000	2500	792	0.00%
2500	3000	59	0.00%
3000	3500	4	0.00%
3500	4000	2	0.00%
4000	4500	1	0.00%
4500	5000	1	0.00%
5000	5500	1	0.00%
5500	6000	1	0.00%
6000	6500	1	0.00%
6500	7000	1	0.00%
7000	7500	1	0.00%
7500	8000	1	0.00%
8000	8500	1	0.00%
8500	9000	1	0.00%
9000	9500	1	0.00%
9500	10000	1	0.00%



Histogram of elevation values for north hemisphere

NOTES ON BASE
This sheet is one in a series of topographic maps that presents color-coded topographic data digitally merged with shaded relief data.
ADOPTED FIGURE
The figure for the Moon, used for the computation of the map projection, is a sphere with a radius of 1737.4 km (Staidman and others, 2001). Because the Moon has no surface water, and hence no sea level, the datum (the 0 km contour) for elevations is defined as the radius of 1737.4 km. Coordinates are based on the mean Earth-polar axis (M.E.P.A.) coordinate system, the z axis is the axis of the Moon's rotation, and the x axis is the mean Earth direction. The center of mass is the origin of the coordinate system (Davies and Colvin, 2000). The equator lies in the x-y plane and the prime meridian lies in the x-z plane with east longitude values being positive.
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The projection is Lambert Azimuthal Equal Area Projection. The scale factor at the central latitude and central longitude point is 1:10,000,000. For the north hemisphere the central latitude and central longitude point is at 90° north and 0°. For the south hemisphere the central latitude and central longitude point is at 90° south and 0°.
CONTROL
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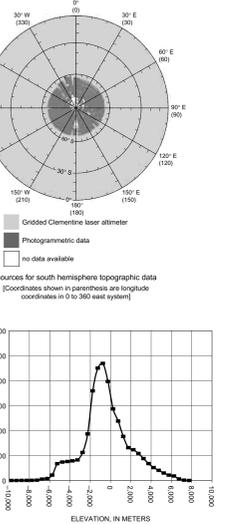
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The shaded relief data were originally published as a series of 1.5 million shaded relief maps. This series included three U.S. Geological Survey maps: I-1218-B, Shaded Relief Map of the Lunar Far Side, 1968; I-1326-A, Shaded Relief Map of the Lunar Near Side, 1981; and I-2276, Sheet 2, Shaded Relief Map of the Lunar Near Side, 1992. These data were digitized and mosaicked into a single digital file. An original control for the shaded relief maps was based on horizontal control used to either the Lunar Positional Reference of 1973 (Schlatterman, 1975) or the Apollo control system of 1973. Positional discrepancies as large as 25 km at a map scale existed in the original shaded relief base. To improve the accuracy, digital shaded relief data were aligned with a mosaic produced from Clementine 750 nm images (Eliason, 1997; Isbell and others, 1997; Eliason and others, 1999). This alignment process consisted of picking features (commonly called match or tie points) that were visible in both the shaded relief data and the Clementine mosaic. To accomplish this, the files were divided into three areas: north pole, equatorial region, and south pole. They were aligned first in the equatorial region and then in the polar regions. Within the equatorial region, an area extending from 60° S. to 60° N., approximately 1100 points were picked. Within the north polar region, an area from 57° N. to 90° N., approximately 1900 points were picked. Within the south polar region, an area from 57° S. to 90° S., approximately 1100 points were used to warp the shaded relief map to match the Clementine mosaic. The Clementine mosaic has a positional accuracy of 500 m (1:6000) (Eliason and others, 1999). Vertical control is based on measurements from the Clementine laser altimeter that collected data between 79° S. and 81° N. The along-track spacing of these measurements varied over time, smooth mare surfaces an along-track spacing of 20 km was achieved, where the instrument lost lock over rough highland terrain, the spacing degraded to 100 km.

control was established by using the global topographic gridded digital terrain model developed from the Clementine laser altimeter points to estimate elevation values for the Clementine match points. To improve the geometry of the control network for the photogrammetric analysis, the Clementine match points, which were selected to tie two images together in order to build the Clementine global mosaic, were transferred to all images that contained the point. Additional points were selected to have four well-distributed points per image, where possible. Analytical aerotriangulation, a weighted least squares process, is used to solve for all the parameters of the photogrammetric project. These parameters include image sensor position and angles; latitude, longitude, and elevation of match points; and image coordinates of match points. Adjusting the weight assigned to a parameter determines whether values with high weight are held to the original estimate or values with low weight are allowed to float and a new value determined for the parameter. The parameters with the most error in their original estimate for their values are the image sensor angles, so they are given a low weight. The latitude and longitude values of Clementine match points are given a high weight to the solution holds to the Clementine global mosaic horizontal coordinates. Weights for the elevation values are varied depending on the horizontal distance to a Clementine laser altimeter point: match points within 2000 m of a Clementine laser altimeter point are given a high weight; match points between 2000 m and 5000 m from a Clementine laser altimeter point are given a medium weight; and match points greater than 5000 m from a Clementine laser altimeter point are given a low weight. This weighting allows the vertical control to be bridged between areas of known vertical control (the area covered by Clementine laser altimeter measurements) and the areas void of control (the area over the poles). The Clementine mission collected both oblique and vertical images over the polar regions; these images from stereo pairs that can be used photogrammetrically to collect topographic data. In the north pole region (90° N. to 64° N. latitude) data were collected from 640 stereo models. Over the south pole region (90° S. to 63° S. latitude) topographic data were collected from 667 stereo models. Topographic data were collected within each stereo model in an elevation point spacing of 1 km in the x and y directions. This spacing resulted in 1,437,368 points being collected in the north pole region and 1,274,872 points being collected in the south pole region. On average, over the area that data were collected (79° S. and 81° N.), the Clementine laser altimeter data collected an elevation value for every 514 km², and the photogrammetric data collected an elevation value for every 1.3 km² in the north pole and 1.2 km² in the south pole. The photogrammetric data were merged and vertically transferred to align with the Clementine laser altimeter data to form the final digital terrain model (Rosick and others, 2001). Merging the topographic data required an iterative process to reduce the error between the photogrammetric data and the topographic data

derived from the Clementine laser altimeter data. At first the photogrammetric topographic data exhibited a systematic stair-step error, in that stereo models closer to the poles had a systematic bias to be higher than stereo models further from the poles. When this bias was removed, the resulting digital terrain model had a bowl-shape appearance with a low spot near the pole. Each stereo model was tilted to remove the bowl shape for stereo models in the north pole area this was 800 m per degree and for stereo models in the south pole area this was 1 km per degree. To fit the stereo models to the Clementine laser altimeter data, an initial set of stereo models that overlaid the Clementine laser altimeter data was adjusted to vertically align with the Clementine laser altimeter data. Stereo models that overlaid the previously adjusted stereo models were adjusted further, this process continued until all stereo models were adjusted to align vertically.
OTHER SOURCES OF LUNAR ELEVATION DATA
Two other recent sources of lunar elevation data are from work done by Dr. Tony Cook of the National Air and Space Museum and by Dr. Jean-Luc Margot while at Cornell University. The elevation data collected by Dr. Cook is derived from overlapping Clementine images from the same orbit. This method provides a larger set of images to use in collecting topographic data. Elevation values that are collected are relative to each other and are not known in an absolute sense relative to some datum. To provide an estimate of the absolute height, the elevation values are matched to the Clementine altimeter data and to elevation values previously collected where Clementine altimeter data does not exist. Traditional photogrammetric solutions employ a least squares method to distribute the error throughout the model, whereas the method employed by Dr. Cook aligns the derived data with the Clementine altimeter data where it exists and in areas where there is no data the relative elevation measurements are aligned with previous aligned data until the area is connected back with Clementine data. The error in unknown areas is unconstrained, but the method does close back to an area with known elevation values, which constrains the error at the starting and ending spots (Cook and others 2000). The elevation data collected by Dr. Margot is derived from radar interferometry using the Goldstone radar. Over the polar region, five images were collected and mosaicked together. To collect elevation information, the radar signal is sent from one antenna and is received by at least two antennas. The phase difference in the received signal from the two radar sites provides information about the elevation. The elevation information derived from radar interferometry is relative and requires the use of the Clementine altimeter points to provide absolute elevation information relative to a datum (Margot and others 1999, 2000).

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Elevation values for the south hemisphere binned in increments of 500 meters. Count is number of elevation points within the bin. Percent is count for indicated elevation values divided by total count.

Minimum	Maximum	Count	Percent
-10000	-9500	222	0.00%
-9500	-9000	543	0.01%
-9000	-8500	1,079	0.01%
-8500	-8000	2,074	0.01%
-8000	-7500	3,179	0.02%
-7500	-7000	6,236	0.03%
-7000	-6500	10,393	0.05%
-6500	-6000	18,548	0.09%
-6000	-5500	32,705	0.16%
-5500	-5000	58,862	0.29%
-5000	-4500	104,019	0.52%
-4500	-4000	191,176	0.95%
-4000	-3500	351,333	1.72%
-3500	-3000	650,490	3.17%
-3000	-2500	1,235,289	6.02%
-2500	-2000	2,308,202	11.35%
-2000	-1500	4,244,140	20.59%
-1500	-1000	7,690,187	37.56%
-1000	-500	13,125,259	63.86%
-500	0	14,222,579	69.00%
0	500	10,022,480	48.99%
500	1000	902,322	4.36%
1000	1500	88,066	0.43%
1500	2000	8,427	0.04%
2000	2500	792	0.00%
2500	3000	59	0.00%
3000	3500	4	0.00%
3500	4000	2	0.00%
4000	4500	1	0.00%
4500	5000	1	0.00%
5000	5500	1	0.00%
5500	6000	1	0.00%
6000	6500	1	0.00%
6500	7000	1	0.00%
7000	7500	1	0.00%
7500	8000	1	0.00%
8000	8500	1	0.00%
8500	9000	1	0.00%
9000	9500	1	0.00%
9500	10000	1	0.00%

Histogram of elevation values for south hemisphere

COLOR-CODED TOPOGRAPHY AND SHADED RELIEF MAP OF THE LUNAR NORTH AND SOUTH HEMISPHERES L 10M x 900 RTK